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temperature will have to take into account the readings at intermediate times.

1894.	Level.	Azimuth.	Temp.	1894.	Level.	Azimuth.	Temp.
	s.	s.	°		s.	s.	°
Oct. 2	—.05	+.19	56	Nov. 5	—.34	+.01	63
3	—.17	+.27	68	6	—.39	—.08	65
4	—.25	+.22	71	7	—.37	.00	67
5	—.29	+.08	67	8	—.40	—.14	65
8	—.28	—.06	64	9	—.43	—.23	64
9	—.29	—.06	63	10	—.42	—.23	67
10	—.33	—.20	58	12	—.43	. . .	65
12	—.32	—.12	60	14	—.44	—.29	71
15	—.30	+.04	71	16	—.39	—.27	55
16	—.30	—.04	64	19	—.42	—.19	61
Southeast Storm.				23	—.43	—.08	56
22	—.16	+.23	53	Southeast Storm.			
24	—.19	+.14	52	28	—.33	—.03	38
25	—.25	+.22	58	30	—.25	+.07	36
29	—.27	+.17	55	. . .	. . .	. . .	. . .
30	—.28	+.14	55	. . .	. . .	. . .	. . .

The collimation correction, as usual, has been very small, and nearly constant. Its value rarely reaches $^{s}.01$, which is but one-seventh of the thickness of the thread whose position is to be determined.

R. H. T.

The Velocity of ζ *Herculis* in the Line of Sight.

The 3d magnitude star ζ *Herculis* is remarkable for its very great velocity in the line of sight. Dr. A. Belopolsky, the able head of the Pulkowa Observatory Spectroscopic Department, took seven photographs of this star's spectrum in 1893, which showed that the star was approaching the solar system at the rate of 70 kilometers (43 miles) per second. Belopolsky's result was confirmed by MM. Deslandres and Millochau of the Paris Observatory, who obtained 62 kilometers per second. This is the greatest velocity in the line of sight yet found for any of the bright stars. The star whose known velocity is next in magnitude is α *Tauri* (*Aldebaran*), for which Vogel and Scheiner obtained a recession of 48.5 kilometers per second. It should be noted, however, that the solar system is approaching ζ *Herculis* and receding from α *Tauri*, so that the observed velocities of both stars must really be decreased by nearly the amount of the Sun's motion. This probably lies between 10 and 30 kilometers per second.

If we except some of the new stars, which seem to have abnormally high velocities, it is probable that the greatest sidereal velocity in the line of light yet observed is that of the planetary nebula G. C. 4373. Professor KEELER in 1890 showed that this object has a velocity of approach of 65 kilometers (40 miles) per second. The nebula is situated very near the pole of the ecliptic, and its relative velocity is therefore unaffected by the motion of the solar system. W. W. C.

THE ORBITAL MOTION OF δ *Cephei.*

The star δ *Cephei* was discovered to be variable by GOODRICKE in 1784. It oscillates between the 3.7 and 4.9 magnitudes in $5^d\ 8^h\ 47^m\ 39^s$, ascending from 4.9 to 3.7 in about 38 hours and descending to 4.9 in the remaining 91 hours. Dr. BELOPOLSKY of Pulkowa has recently observed the velocity of this star in the line of sight. He has found that the velocity varies in a period coincident with the period of the variation in brightness, and that the star is moving in a very excentric orbit about a dark or relatively dark companion. Thus the variation in brightness is presumably caused by one star eclipsing a part of the other star.

W. W. C.

THE MERIDIAN CIRCLE IN JANUARY.

The following list of corrections, determined on regular observing nights, shows the good performance of the instrument at temperatures comparatively low for Mount Hamilton :

1895.	c.	b.	a.	Zenith.	Temp.
January 25.	.00	− .50	+ .08	20″.0	37°
26.	.00	− .51	+ .05	20 .3	34
28.	+ .01	− .53	− .03	21 .2	32
29.	.00	− .54	+ .01	19 .7	44
30.	− .01	− .57	− .01	19 .9	44
31.	.00	− .59	− .11	20 .1	45
February 1.	.00	− .60	− .10	20 .2	47

There is rather large variation in the reading of the zenith point, showing some dependence upon the temperature. Ad-